

that this radiation is almost entirely absorbed in traversing a grain of AgBr ($<1\mu$). This seems unlikely. A recent alternative explanation due to Schönberg¹⁶ is that most of the relativistic increase arises from ionization and excitation.

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* Maurice M. Shapiro and Bertram Stiller, *Bull. Am. Phys. Soc.* **27**, No. 4, p. 18.

¹ G. P. S. Occhialini, *Nuovo cimento* **6**, 413 (1949); also see discussion by P. M. S. Blackett, *Nuovo cimento* **6**, 427 (1949).

² P. H. Fowler, *Phil. Mag.* **41**, 169 (1950). The asymptotic zero slopes in Fig. 3 imply that no relativistic rise was observed.

³ E. Pickup and L. Voyvodic, *Phys. Rev.* **80**, 89 (1950).

⁴ H. Messel and D. M. Ritson, *Phil. Mag.* **41**, 1129 (1950). These authors correctly stress one source of confusion: the failure to distinguish between total ionization loss and the limited loss along a track measured by Ag grain density. By "ionization" we shall mean the latter loss, arising from energy transfers to electrons up to, say, 5 kev.

⁵ Bradt, Kaplon, and Peters, *Helv. Phys. Acta* **23**, 24 (1950).

⁶ I. B. McDiarmid, *Phys. Rev.* **84**, 851 (1951).

⁷ A. H. Morrish, *Phil. Mag.* **43**, 533 (1952).

⁸ M. Schönberg, *Nuovo cimento* **8**, 159 (1951).

⁹ E. Fermi, *Phys. Rev.* **57**, 485 (1940).

¹⁰ O. Halpern and H. Hall, *Phys. Rev.* **73**, 477 (1948).

¹¹ References to the literature on ionization theory will be found in the monograph by A. Bohr, *Det. Kgl. Danske Videnskab. Selskab, Mat.-fys. Medd.* **24**, No. 19 (1948).

¹² The plates were processed as described by Stiller, Shapiro, and O'Dell, [*Phys. Rev.* **85**, 712 (1952); paper to appear in *Rev. Sci. Instr.*], so as to give a favorable ratio of g to background, hence the low values of g ($g_{\min} = 19.0$ grains/100 microns) compared to those attainable by stronger development. Ag "blobs" were infrequent since g was low. Grains were counted individually, not by blobs. Through 95 percent of the emulsion depth, g was remarkably uniform. Counting was avoided in the top 5 percent.

¹³ Camerini, Lock, and Perkins, *Progress in Cosmic-Ray Physics* (North-Holland Publishing Company, Amsterdam, 1952), p. 24.

¹⁴ E. Fermi, *Prog. Theoret. Phys.* **5**, 570 (1950).

¹⁵ L. Voyvodic, Report, Bristol Conf. on Heavy Mesons, Dec., 1951, p. 16.

¹⁶ M. Schönberg, *Nuovo cimento* **9**, 372 (1952).

The Decay of Kr⁸⁷

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THE decay of 78-min Kr⁸⁷ has not previously been studied β -spectrometrically. By absorption measurements on electromagnetically separated Kr⁸⁷, formed in fission, Koch *et al.*¹ found a maximum energy of the β -particles of 3.2 Mev.

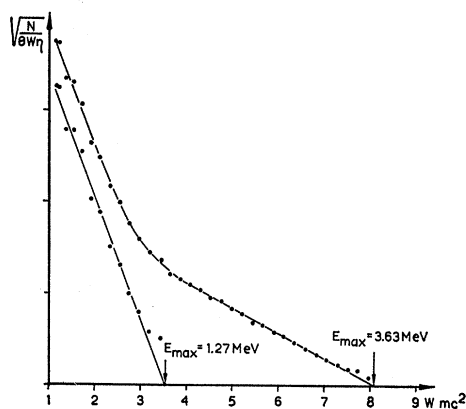


FIG. 1. Fermi plot of the β -continuum of Kr⁸⁷.

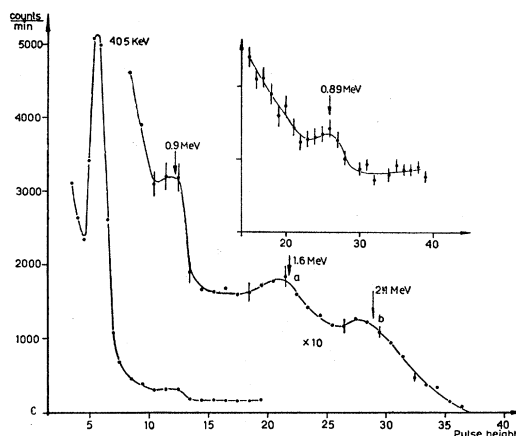


FIG. 2. Scintillation spectrum of Kr⁸⁷. The inset shows a separate measurement of the 0.89-Mev line.

In the present investigation Kr⁸⁷ has been electromagnetically separated from other rare gas isotopes formed in uranium fission. The isotope of mass number 87 was collected on an Al foil of thickness 0.15 mg/cm². This foil was then used as a β -source in a magnetic lens spectrometer (resolution ~ 6 percent).

The β -continuum appears to be complex. The Fermi plot (Fig. 1) shows the presence of two β -components of energies 3.63 ± 0.07 and 1.27 ± 0.1 Mev, with relative intensities 75 percent and 25 percent, respectively. The values of $\log ft$ are found to be 7.3 (the Fermi plot having "not forbidden" form) and 5.9, respectively.

The scintillation spectrum of Kr⁸⁷ (Fig. 2) shows the presence of a pronounced photopeak of a γ -ray of energy 405 ± 20 kev. The

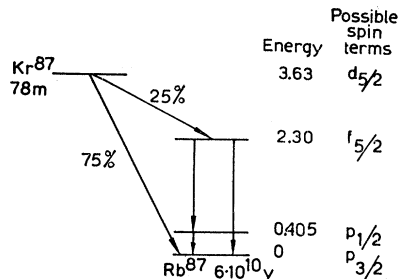


FIG. 3. Tentative decay scheme for Kr⁸⁷.

weak line at 0.89 Mev may be interpreted as the pair line of a γ -ray of energy 1.89 ± 0.07 Mev, the Compton peak of which appears at a (Fig. 2). The high energy peak b is most probably due to Compton quanta corresponding to a γ -ray of energy ~ 2.3 Mev. The scintillation spectrum ends at 2.6 Mev.

The data thus obtained can be accounted for by the decay scheme of Fig. 3. [Spin terms are given according to the strong spin orbit coupling theory.²] If this interpretation (Fig. 3) is correct, it is worth noting that the $f_{5/2}$ level is so far above the already filled $f_{7/2}$ level.

¹ Koch, Kofeod-Hansen, Kristensen, and Drost-Hansen, *Phys. Rev.* **76**, 279 (1949).

² Haxel, Jensen, and Suess, *Phys. Rev.* **75**, 1766 (1949); M. G. Mayer, *Phys. Rev.* **75**, 1969 (1949).